

AMBER/VLTI high spectral resolution observations of the Br γ emitting region in HD 98922

A compact disc wind launched from the inner disc region \star

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ABSTRACT

Context. High angular and spectral resolution observations can provide us with fundamental clues to the complex circumstellar structure of young stellar objects (YSOs) and to the physical processes taking place close to these sources.

Aims. We analyse the main physical parameters and the circumstellar environment of the young Herbig Be star HD 98922.

Methods. We present AMBER/VLTI high spectral resolution ($R=12\,000$) interferometric observations across the Br γ line, accompanied by UVES high-resolution spectroscopy and SINFONI-AO assisted near-infrared (NIR) integral field spectroscopic data. To interpret our observations, we develop a magneto-centrifugally driven disc-wind model.

Results. Our analysis of the UVES spectrum shows that HD 98922 is a young ($\sim 5 \times 10^5$ yr) Herbig Be star (SpT=B9V), located at a distance of 440 ± 60 pc, with a mass accretion rate (\dot{M}_{acc}) of $\sim (9 \pm 3) \times 10^{-7} M_{\odot} \text{ yr}^{-1}$. SINFONI K -band AO-assisted imaging shows a spatially resolved circumstellar disc-like region (~ 140 AU in diameter) with asymmetric brightness distribution. Our AMBER/VLTI UT observations indicate that the Br γ emitting region (ring-fit radius $\sim 0.31 \pm 0.04$ AU) is smaller than the continuum emitting region (inner dust radius $\sim 0.7 \pm 0.2$ AU), showing significant non-zero V-shaped differential phases (i.e. non S-shaped, as expected for a rotating disc). The value of the continuum-corrected pure Br γ line visibility at the longest baseline (89 m) is $\sim 0.8 \pm 0.1$, i.e. the Br γ emitting region is partially resolved. Our modelling suggests that the observed Br γ line-emitting region mainly originates from a disc wind with a half opening angle of 30° , and with a mass-loss rate (\dot{M}_w) of $\sim 2 \times 10^{-7} M_{\odot} \text{ yr}^{-1}$. The observed V-shaped differential phases are reliably reproduced by combining a simple asymmetric continuum disc model with our Br γ disc-wind model.

Conclusions. In conclusion, the Br γ emission of HD 98922 can be modelled with a disc wind that is able to approximately reproduce all interferometric observations if we assume that the intensity distribution of the dust continuum disc is asymmetric.

Key words. stars: formation – stars: circumstellar matter – stars: pre-main sequence – stars: variables: Herbig Ae/Be: individual objects: HD 98922 – techniques: interferometric – techniques: high angular resolution

1. Introduction

Herbig Ae/Be stars are intermediate-mass pre-main-sequence stars ($2\text{--}10 M_{\odot}$; see e.g. Herbig 1960; Waters & Waelkens 1998).

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\star Based on observations collected at the VLT (ESO Paranal, Chile) with programmes 075.C-0637(A), 083.C-0236(A-D), 090.C-0192(A), 090.C-0378(A) and 090.C-0371(A)

They are the intermediate-mass counterparts of Classical T Tauri stars (CTTSs) and are a fundamental link between low- and high-mass star formation. As is the case for their low-mass counterparts, the circumstellar environment of Herbig Ae/Be stars is not fully understood. In particular, their innermost disc regions, within 1 AU from the central sources, harbour both accretion and ejection processes, whose study is of fundamental relevance to understanding how stars form. Herbig Ae/Be stars display several features indicative of their complex circumstellar environments and of the accretion and ejection processes taking place close to the sources: infrared excess indicating circumstellar material and disc, UV excess and veiling (Donehew & Brittain 2011) indicative of accretion, as well as several permitted and forbidden lines which can be used to trace both accretion and ejection processes (e.g. H I, He I, or [O I], [Fe II] lines).

However, owing to the small spatial scales involved (1 AU corresponds to 10 milliarseconds -mas- at 100 pc) high angular resolution is required to spatially resolve the inner disc region. Additionally, to study the gas kinematics and disentangle the different processes, high spectral resolution is needed as well. Near-IR interferometry at medium and high spectral resolution has thus become a fundamental tool in probing the accretion/ejection processes in the inner disc regions (see e.g. Eisner et al. 2007; Kraus et al. 2008; Eisner et al. 2009, 2014; Weigelt et al. 2011).

Observations have been mostly focused on the Br γ line (2.166 μ m), which is bright in Herbig Ae/Be stars and can often be associated with accretion or ejection processes. Notably, only five Herbig Ae/Be stars have been observed so far through high spectral resolution NIR interferometry ($R \sim 12\,000$; see Weigelt et al. 2011; Kraus et al. 2012; Garcia et al. 2013; Garcia Lopez et al. 2015; Ellerbroek et al. 2015). In particular, VLTI/AMBER observations of the Br γ at high spectral resolution have been proved effective in resolving the disc-wind region and in inferring the flux contribution of the unresolved magnetosphere or X-wind region to the total line flux. Our group has successfully developed disc-wind, X-wind, and magnetospheric models that match these observations very well (see Grinin & Tambovtseva 2011; Weigelt et al. 2011; Tambovtseva et al. 2014; Garcia Lopez et al. 2015). In this paper we study the circumstellar environment of HD 98922, through VLTI/AMBER interferometry, along with UVES (high-resolution UVB spectroscopy) and SINFONI-AO assisted (NIR integral field spectroscopy) ancillary data.

HD 98922 is a Herbig Be star of spectral-type B9V (Houk 1978), although a more recent analysis suggests it might be a later spectral type (A2III; Hales et al. 2014). Its SED has a large infrared excess, indicating the presence of an extended dusty disc (Malfait et al. 1998). Several optical and near-IR emission lines prove its strong circumstellar activity. In particular, from the analysis of the [O I] emission line profiles, Acke et al. (2005) conclude that this line cannot originate from the self-shadowed dusty disc surface, but likely from a rotating gaseous disc inside the dust-sublimation radius. From the analysis of the Br γ line, Garcia Lopez et al. (2006) derive a high mass-accretion rate ($\dot{M}_{acc} \sim 2 \times 10^{-6} M_{\odot} \text{ yr}^{-1}$, assuming a distance of 540 pc), indicating that the Herbig Be star is still actively accreting. AMBER-LR and MR interferometric studies from Kraus et al. (2008) show that the Br γ line visibility increases with respect to the continuum visibility. In particular, the Br γ line-emitting region is not resolved with baseline lengths up to ~ 60 m (an upper limit of 1 mas, or 0.5 AU at a distance of 540 pc, is reported), whereas the continuum-emitting region ring-fit diameter is estimated to be 4.6 ± 0.1 mas (or 2.48 ± 0.05 AU at a distance of 540 pc). The authors conclude that the Br γ line could originate

from a stellar wind, X-wind, or magnetospheric accretion, and they favour the last scenario. Moreover, a strong wind/outflow activity is also observed from the P Cygni profiles of the H α and Si II lines ($v_{rad} \sim 300 \text{ km s}^{-1}$; Grady et al. 1996) at optical wavelengths, as well as from the He I line profile (Oudmaijer et al. 2011) in the near-IR.

Visual extinction values towards the object are between 0.3 and 0.5 mag, and the average magnitude in the V band is 6.67 mag (see e.g. Manoj et al. 2006). The inclination of the HD 98922 system axis with respect to the plane of the sky is poorly known. A very rough estimate of 45° was derived in Blondel & Djie (2006) by modelling the UV spectrum, whereas Hales et al. (2014) report an estimate of 20° from CO modelling.

Finally, the distance to the source is quite uncertain. Trigonometric distance measured with the HIPPARCOS satellite (van den Ancker et al. 1998) gives a value of $d > 540$ pc, which is well beyond the HIPPARCOS observational range ($d \sim 350\text{--}400$ pc). A later revision of these data provided a larger value of $d = 1190^{+930}_{-390}$ (van Leeuwen 2007). This value translates into an extremely large stellar luminosity for HD 98922 of $\sim 10^4 L_{\odot}$. As a consequence its position in the Hertzsprung–Russell diagram is well outside the locus of the pre-main sequence Herbig Ae/Be stars (see e.g. Fig. 4 in Alecian et al. 2013, see also a more detailed discussion in Sect. 3.1.1), namely its position cannot be reproduced by using the evolutionary tracks of Herbig Ae/Be stars. Recently, the analysis of Hales et al. (2014) has provided a closer distance of 507^{+131}_{-104} pc. It is also worth noting that this source is a wide binary ($\sim 8''$, $PA = 343^\circ$ Dommangeat & Nys 1994) and it is possible that a second closer companion might also be present, as hinted by spectro-astrometric observations ($< 0''.5$, $PA \sim 0^\circ$ Baines et al. 2006).

The paper is structured as follows. Spectroscopic and interferometric observations and data reduction are presented in Sect. 2. Section 3 reports results from UVES high-resolution spectroscopy, along with the derived stellar physical parameters and distance. Section 4 presents the results of our SINFONI observations. We present our interferometric results across the Br γ line, along with its geometrical and physical modelling in Sects. 5 and 6. Section 7 provides a discussion on the origin of the Br γ emitting region and of the asymmetries observed in the interferometric observables. Finally, our conclusions are presented in Sect. 8.

2. Observations and data reduction

2.1. VLT/UVES spectroscopy

HD 98922 was observed on 21 March 2005 with UVES (Dekker et al. 2000), the echelle spectrograph mounted on the VLT/UT2. The observations were performed with the blue arm, covering a spectral range between $\sim 3750 \text{ \AA}$ and $\sim 4990 \text{ \AA}$. A slit width of $0''.4$, which gives a spectral resolution of 80 000, was adopted, and the total integration time was 100 s. Wavelength calibration frames were taken with a long slit and a ThAr arc lamp. The data were reduced using the UVES pipeline v3.2.0 (Ballester et al. 2000) available from the ESO Common Pipeline Library. We converted the wavelength scale to the heliocentric rest frame. No standard star is available to provide an accurate flux calibration. Owing to the very small photometric variability of HD 98922 (de Winter et al. 2001), we adopt the photometric value reported in Manoj et al. (2006) ($m_B = 6.81$ mag) to calibrate the spectrum.

2.2. SINFONI AO spectral-imaging

HD 98922 was observed on 2 February 2013 with the VLT/SINFONI integral field spectrograph (Eisenhauer et al. 2003) in the K-band ($1.95\text{--}2.45\mu\text{m}$) at the highest spectral ($R \sim 4000$) and spatial resolution (12.5×25 mas pixel scale). The detector integration time (DIT) and number of sub-exposures per frame were 2 s and five, respectively, both for the on-source (target) and the off-source (sky) position. In total, eight target-sky pairs were obtained at position angles (PAs) 0° , 90° , 180° , and 270° (two pairs per PA), corresponding to a total integration time on-source of 80 s. These observations were AO-assisted (using HD 98922 as a natural guide star), covering a field of view (FoV) of $0.8'' \times 0.8''$.

To correct for atmospheric transmission, observations of a telluric standard star of spectral type A were also performed. The main data reduction process was done using the SINFONI data-reduction pipeline in GASGANO (Modigliani et al. 2007), i.e. dark and bad pixel masks, flat-field corrections, optical distortion correction, and wavelength calibration using arc lamps. As a result, we obtained four calibrated datacubes, one for each observed PA.

We then measured the full width at half maximum (FWHM) of both target and standard star on the final datacubes to estimate the achieved spatial resolution, which is ~ 60 mas. From each datacube the stellar spectrum was then extracted, and continuum and Bry line images were created by collapsing the datacubes along the z-axis, from 2.15 to $2.16\mu\text{m}$ and from 2.164 to $2.168\mu\text{m}$, respectively.

2.3. VLTI/AMBER/FINITO Interferometry

The log of our interferometric observations is reported in Table 1. HD 98922 was observed with AMBER (Petrov et al. 2007), the NIR beam-combiner of the Very Large Telescope Interferometer (VLTI) operated by ESO, during two different runs on 26 December 2012 and 28 February 2013. On both occasions we employed AMBER's high spectral resolution mode in the K-band (HR mode; $R = 12\,000$ or $\Delta v \sim 25\text{ km s}^{-1}$) covering the spectral range from 2.147 to $2.194\mu\text{m}$, centred on the Bry line emission (at $2.166\mu\text{m}$). Observations were conducted with the UT2-UT3-UT4 telescopes, using the fringe tracker FINITO (Gai et al. 2004) for co-phasing and a detector integration time of 1.0 s per interferogram. In the first run 100 interferograms were recorded, whereas ~ 1500 were recorded in our second run. Stars HD 60276 and HD 103125 were observed with the same observational settings and were used as interferometric calibrators to derive the transfer function for the first and second run, respectively. The calibration datasets for the first and second runs consist of 200 and 980 calibrator interferograms, respectively.

To reduce our interferograms, we used our own data reduction software based on the P2VM algorithm (Tatulli et al. 2007), which provides us with wavelength-dependent visibilities, wavelength-differential phases, closure phases, and wavelength-calibrated spectra.

The fringe tracking performance of FINITO typically differs during the observations of target and calibrator. Therefore, we used the archival low-resolution visibilities of HD 98922 (216 different visibility uv points, mostly taken during the AMBER's GTO) to calibrate the continuum visibilities at the six uv points of our HR observations. This calibration method provides us with a visibility error of about $\pm 5\%$ due to the large number of employed LR visibility measurements. By using low-resolution

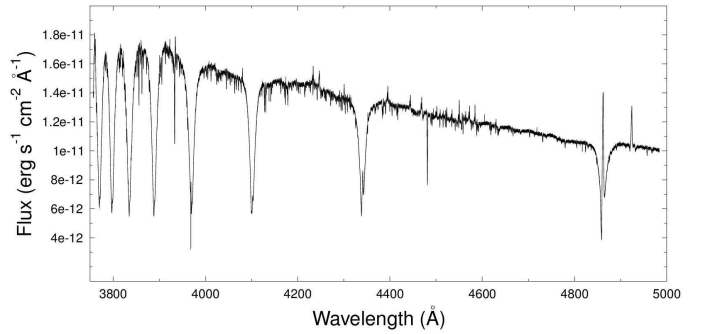


Fig. 1. UVES/VLT flux-calibrated spectrum of HD 98922.

spectroscopic data taken at different epochs to calibrate our data, we are implicitly assuming that both size and morphology of the object have not changed more than the data uncertainties allow.

The wavelength calibration of the AMBER data was refined using the numerous telluric lines present in the $2.15\text{--}2.19\mu\text{m}$ region (see Weigelt et al. 2011, for more details on the wavelength calibration method). We estimate an uncertainty in the wavelength calibration of $\sim 0.2\text{ Å}$ or $\sim 3\text{ km s}^{-1}$.

3. Results from UVES spectroscopy

3.1. Line recognition

The UVES flux-calibrated spectrum of HD 98922 is presented in Figure 1 and shows the descending stellar continuum with several strong absorption lines and few faint emission features. The same spectrum, normalised to the continuum, is shown in Figure 2 (black line). The most prominent features, both in absorption and emission, are also labelled in Fig. 2.

The most prominent absorption features are the H I lines from the Balmer series (from H 11 to H β), which partially show emission due to the circumstellar environment. These lines show P Cygni profiles (see e.g. the H β line vs. its theoretical profile in Fig. 2: the peak in emission is blue-shifted with respect to the zero velocity, and the blue-shifted absorption is deeper than the theoretical value), indicative of winds or outflows. The remaining absorption features are mostly photospheric from atomic lines like He I, Ca II, Fe I, Fe II, Mg I, Mg II, Si I, and Si II. Other emission features, mostly Fe II and Ti II lines, located between ~ 4200 and 4600 Å , arise from circumstellar activity. These lines are characterised by a double-peaked profile, are likely produced by self-absorption of the photospheric feature, and have a broad pedestal (full width zero intensity, $FWZI \sim 4\text{--}4.4\text{ Å}$, i.e. $\Delta v \sim 270\text{--}290\text{ km s}^{-1}$), which originates in a region with a large radial velocity structure, likely located in the magnetospheric accretion funnels or in the hot inner region of the circumstellar disc (Beristain et al. 1998; Aspin et al. 2010). The strongest Fe II emission line (at 4925 Å) shows a P Cygni profile.

From the analysis of the H 11 line, we derive a value of $\sim 8\text{ km s}^{-1}$ for the stellar radial velocity (v_{rad}) with respect to the local standard of rest. Notably, the observed spectral lines do not show any indication of binarity: no splitting and no additional absorption lines are detected in the spectrum.

3.1.1. Stellar physical parameters and distance

To allow both identification and analysis of the observed line profiles we used synthetic spectra from the PHOENIX

Table 1. Log of the VLTI/AMBER/FINITO observations of HD 98922 and calibrators.

HD 98922 Observation Date	Time [UT]		Unit Telescope array	Spectral mode ^c	Wavelength range (μm)	DIT ^a (s)	N ^b	Seeing ($''$)	Baselines (m)	PAs ($^\circ$)
2012 Dec 26	09:02	09:20	UT2-UT3-UT4	HR-K-F	2.147-2.194	1	100	0.7-0.8	44.1/60.6/89.4	219.3/103/76.7
2013 Feb 28	04:50	05:32	UT2-UT3-UT4	HR-K-F	2.147-2.194	1	1540	0.9-1.2	43.9/61/89.4	220.9/104.9/78.8

Calibrator Name	Date	Time [UT]		UT array	Spectral mode ^c	Wavelength range (μm)	DIT ^a (s)	N ^b	Seeing ($''$)	Uniform-disc diameter ^d (mas)
HD 60276	2012 Dec 26	07:50	08:03	UT2-UT3-UT4	HR-K-F	2.147-2.194	1	200	0.9-1.0	0.92 \pm 0.06
HD 103125	2013 Feb 28	04:08	04:42	UT2-UT3-UT4	HR-K-F	2.147-2.194	1	70	0.9-1.5	0.88 \pm 0.06

Notes. ^(a) Detector integration time per interferogram. ^(b) Number of interferograms. ^(c) High spectral resolution mode in the *K*-band using the fringe tracker FINITO. ^(d) UD diameter taken from JMMC Stellar Diameters Catalogue - JSDC (Lafrasse et al. 2010).

library (Husser et al. 2013), which provides high-resolution ($R=500\,000$) synthetic spectra based on the stellar atmosphere code PHOENIX. The spectral resolution was then reduced to that of the UVES observations by convolving the synthetic spectra with a Gaussian profile (IRAF task *gauss*). We constructed a grid of spectra and varied different stellar parameters, namely T_{eff} (from 9000 to 11 000 K, $\Delta T = 200$ K), $\log g$ (from 3.0 to 4.5, $\Delta \log g = 0.5$), and $[\text{Fe}/\text{H}]$ (from -1 to 1, $\Delta[\text{Fe}/\text{H}] = 0.5$). Both observed and synthetic spectra were then normalised to the continuum. Our own IDL programme was used to match the UVES spectrum with the best-fitting model in three steps. First, by comparing the wings of the Balmer lines the stellar gravity was inferred. In this case, Balmer lines from H11 to H γ were used (the H β line was excluded because it is too affected by circumstellar emission). Second, the effective temperature (T_{eff}) of the star was obtained by comparing observed and theoretical equivalent widths (EWs) of lines sensitive to T_{eff} , namely the He I lines (at 4026, 4120, 4145, 4470 Å), and the Ti II line at 4176 Å. Finally, the metal abundance was inferred by comparing the iron abundances for a well-defined set of Fe I and Fe II lines. As a result, we obtained the stellar effective temperature, gravity, and metallicity, namely $T_{\text{eff}} = 10\,400 \pm 200$ K (i.e. spectral type B9V), $\log g = 3.5 \pm 0.2$, and $[\text{Fe}/\text{H}] = -0.5 \pm 0.2$. The best-fitting synthetic spectrum, normalised to the stellar continuum, is shown in Figure 2 (red solid line) superimposed on the observed spectrum (in black). For comparison, Fig. 2 also shows a normalised synthetic spectrum (green dashed line) with the stellar parameters derived by Hales et al. (2014), namely $T_{\text{eff}} = 9000$ K (SpT=A2 III), $\log g = 3.0$, and $[\text{Fe}/\text{H}] = -0.5$.

To derive the remaining stellar parameters, we follow Montesinos et al. (2009). We place the estimated values of the stellar parameters on a $\log T_{\text{eff}}$, $\log g$ HR diagram, using evolutionary tracks and isochrones from Siess et al. (2000) with $[\text{Fe}/\text{H}] = -0.5$. This provides us with an estimate of the stellar mass ($M_* = 5.2 \pm 0.2 M_\odot$) and age ($\sim 5 \times 10^5$ yr) for HD 98922, and, in turn, stellar luminosity ($L_* = 640 \pm_{160}^{130} L_\odot$), radius ($R_* = 7.6 \pm_{2}^{0.6} R_\odot$), and visual absolute magnitude ($M_V = -1.75 \pm 0.3$ mag) can be inferred from a ($\log T_{\text{eff}}$ vs. $\log L_*$) HR diagram.

Following Montesinos et al. (2009), we also aim at deriving a better estimate of the distance (d) to HD 98922, from the distance module equation:

$$d = 10^{(m_V - A_V - M_V + 5)/5}. \quad (1)$$

Table 2. HD 98922 stellar and disc parameters

Stellar Parameter	Value	Reference
Distance	$440 \pm_{50}^{60}$ pc	1
T_{eff}	$10\,400 \pm 200$ K	1,2
<i>SpType</i>	B9V	1,2
$\log g$	3.5 ± 0.2	1
$[\text{Fe}/\text{H}]$	-0.5 ± 0.2	1
M_*	$5.2 \pm 0.2 M_\odot$	1
R_*	$7.6 \pm_{2}^{0.6} R_\odot$	1
L_*	$640 \pm_{160}^{130} L_\odot$	1
age	5×10^5 yr	1
M_V	-1.75 ± 0.3 mag	1
m_V	6.77 mag	3
A_V	0.3	3
m_K	4.28	4
$F_{\text{obs}}(K)/F_*(K)$	4.3	1
i	20°	5
$v_{\text{rot}} \sin i$	$50 \pm 3 \text{ km s}^{-1}$	6
\dot{M}_{acc}	$(9 \pm 3) \times 10^{-7} M_\odot \text{ yr}^{-1}$	4 + 1

References. (1) This work; (2) Thé et al. (1994); (3) Manoj et al. (2006); (4) Garcia Lopez et al. (2006); (5) Hales et al. (2014); (6) Alecian et al. (2013);

We assume $m_V = 6.77$ mag and $A_V = 0.3$ (Garcia Lopez et al. 2006), and we obtain $d = 440 \pm_{50}^{60}$ pc. By adopting this distance value, an estimate of the mass accretion rate (\dot{M}_{acc}) can be also inferred from the EW of the Bry line after correcting for the visual extinction and self-absorption from the photosphere (see e.g. Garcia Lopez et al. 2006). By using their published data we obtain $\dot{M}_{\text{acc}} \sim (9 \pm 3) \times 10^{-7} M_\odot \text{ yr}^{-1}$.

All stellar and disc parameters used in this paper, derived in this work or adopted from the literature, are listed in Table 2.

4. Results from SINFONI-AO spectral imaging

Our SINFONI K-band spectral images resolve a bright asymmetric scattered emission (diameter ~ 320 mas at 3σ , or ~ 140 AU at a distance of 440 pc) that extends beyond the symmetric point spread function (PSF; FWHM=60 mas) of HD 98922 (see the upper panel of Figure 3 and its contour levels). This structure is not an artefact and because it is visible in all the spectral images

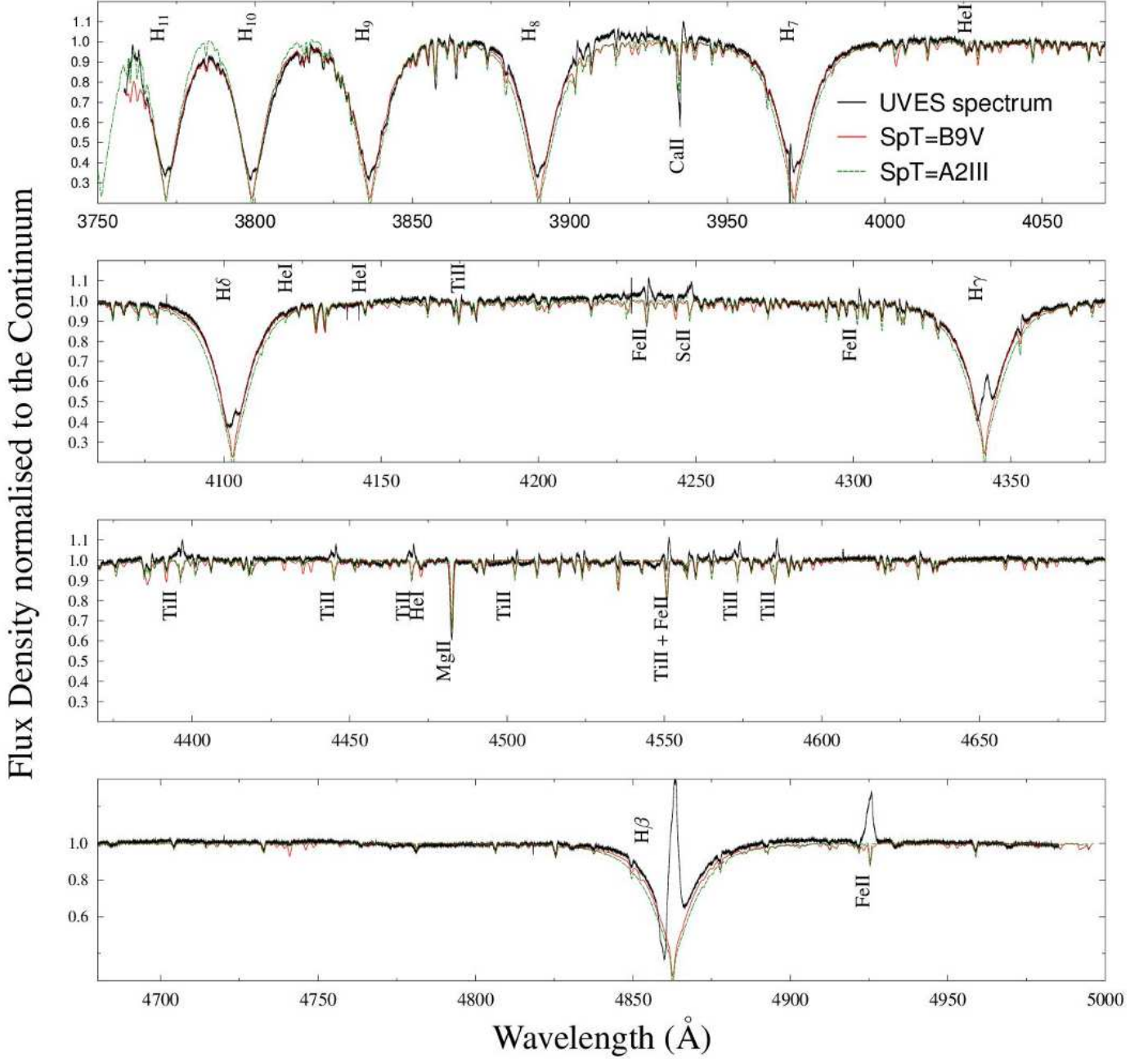


Fig. 2. UVES/VLT high-resolution spectrum of HD 98922 normalised to the continuum (black solid line). The red solid line shows the best-fitting synthetic spectrum, $SpT=B9V$, $\log g=3.5$, $[Fe/H]=-0.5$. For comparison, a synthetic spectrum with parameters $SpT=A2III$, $\log g=3.0$, $[Fe/H]=-0.5$ (green dashed line) is also shown. The most prominent spectral features are labelled.

and is co-rotating in each image FoV according to the different observed PAs.

The bright circumstellar structure most likely originates from the scattered light of the disc. Of particular interest is its asymmetric morphology. The brightness distribution is roughly represented by a symmetric PSF plus a bright elliptical arc towards the north-north-east (NNE) (i.e. a horseshoe-like region), extending from $PA \sim 290^\circ$ to $\sim 110^\circ$ (see upper panel of Fig. 3) and located at ~ 90 mas (or ~ 40 AU) from the source. This region is about 2.5 times brighter than its symmetric counterpart located towards south-south-west (SSW). The value of this ratio was obtained after integrating the flux in two regions (positioned outside the

PSF, i.e. 60 mas), which have the same area and are symmetrical with respect to the central PSF. It is also worth noting that the brightness distribution of the NNE arc is not homogeneous; it shows an elongated peak between $PA \sim 23^\circ$ and $PA \sim 65^\circ$, and it is up to three times brighter than the whole arc-shaped feature. The total flux of this area is about one tenth of the flux of the central PSF. Because of its elongated arc-shape, which cannot be fitted with the instrumental PSF, we exclude that this structure can be the unresolved PSF of a possible companion. Remarkably, our SINFONI-AO continuum images do not show any companion between ~ 800 mas (or ~ 350 AU) and ~ 60 mas (or 26.4 AU) from the central source. The lower limit of the flux

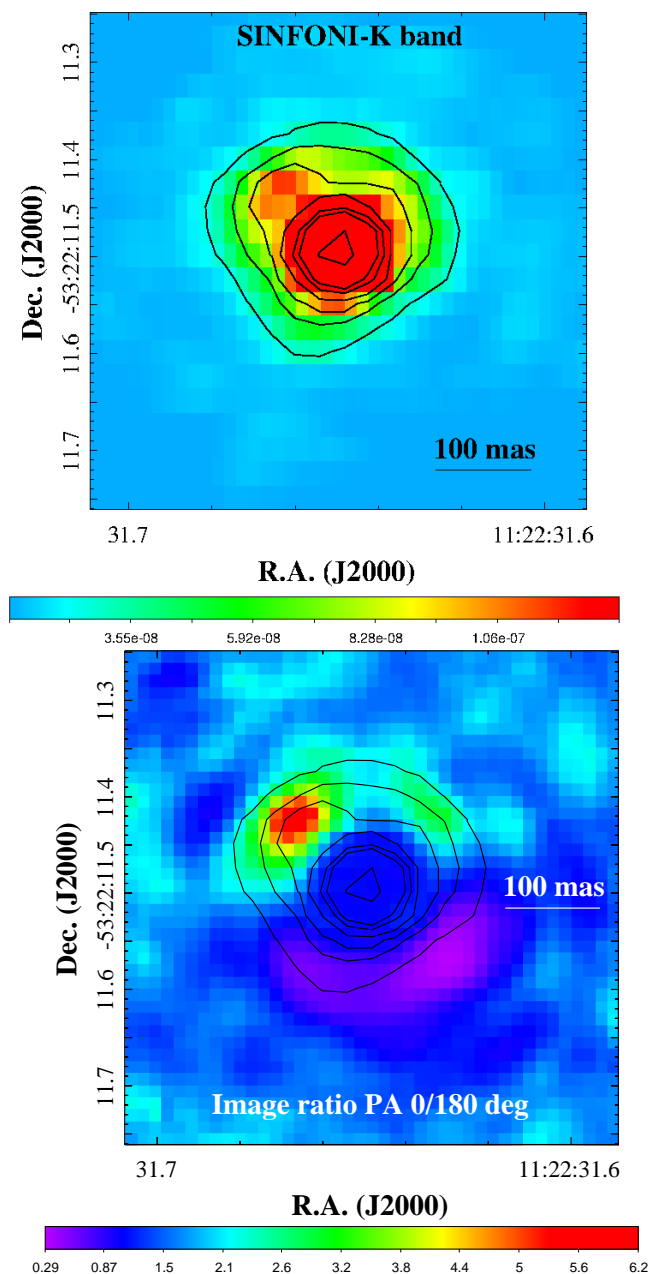


Fig. 3. *Upper Panel:* SINFONI K-band continuum image (between 2.15 and 2.16 μm). Labelled fluxes are in $\text{W m}^{-2} \mu\text{m arcsec}^{-2}$. Contour levels are 4, 5, 50, 150, 250, 400, 500, and 1000 σ . *Lower Panel:* Intensity ratio image of two SINFONI K-band continuum images (between 2.15 and 2.16 μm) taken at PA of 0° and 180°, showing the asymmetric disc emission. Contours show the K-band continuum of the PA=0° image (see upper panel), and give the position of both symmetric PSF in the centre, and asymmetric disc emission. Contour levels are as in the continuum image.

density in our data is $\sim 5 \times 10^{-11} \text{ erg s}^{-1} \text{ cm}^{-2} \mu\text{m}^{-1}$, which corresponds to $K \sim 9.8 \text{ mag}$ or $M_K \sim 1.6 \text{ mag}$ at the estimated distance of 440 pc. Adopting the Siess et al. (2000) evolutionary tracks and assuming an age of $\sim 5 \times 10^5 \text{ yr}$ for any possible undetected companion, the M_K estimate provides us with an upper

limit constraint on the spectral type and mass of SpT $\sim \text{M0}$ and $M_* \sim 0.5 M_\odot$.

To better display the observed brightness asymmetry in the disc, the lower panel of Figure 3 shows an intensity ratio image that is computed from two SINFONI K-band continuum images (between 2.15 and 2.16 μm) with PAs of 0° and 180°, respectively. The central unresolved PSF disappears (ratio value ~ 1), whereas the bright NNE arc becomes clearly visible (ratio values between 1.5 and 5) in the image. Apart from the aforementioned elongated peak, there is a dip at PA $\sim 0^\circ$ and a second arc-shaped peak towards the north-west that is less bright.

It is worth mentioning that the PIONIER/VLTI reconstructed images of the HD 98922 inner disc (see Figure 2 in Kluska et al. 2014), show a similar brightness disc asymmetry, namely the north-eastern side of the disc is brighter than the south-western side. These interferometric data were acquired between 20 December 2012 and 20 February 2013 (Kluska priv. comm.), simultaneous with our AMBER and SINFONI observations. The observed PIONIER disc has, however, a size of $\sim 3\text{--}4 \text{ mas}$ ($\sim 1\text{--}1.5 \text{ AU}$), which is well within the (spatially unresolved) PSF. This suggests that the disc might have a similar morphology from ~ 1 to 60–70 AU from the source and that the origin of this brightness asymmetry might be the same.

The Bry line is the only emission line detected over the K-band continuum of the analysed SINFONI spectral-images. This emission is not spatially resolved, being enclosed within 60 mas (or $\sim 26 \text{ AU}$) from the central source.

5. Results from VLTI/AMBER interferometry and geometric modelling

5.1. Interferometric observables: Visibilities, differential phases, and closure phase

Our VLTI/AMBER observations provide us with four direct observables, namely the Bry line profile, visibilities, differential phases, and closure phases. These observables allow us to retrieve information about the inner-disc emitting region of HD 98922, in particular the size and kinematics of the Bry emitting region, as well as its displacement with respect to the continuum emission.

First of all, we stress that observations in both VLTI runs were taken with the same AMBER settings and UT configurations, which result in similar projected baselines and position angles (see Cols. 10 and 11 of Table 1). The results from both runs are identical inside the error bars. However, those from the second run (28 February 2013) are clearly much less noisy than those from the first run (26 December 2012) owing to the longer total integration time. For this reason, in this section we only present the results from the second run (Figure 4), whereas those from the first run are shown in Appendix A (Figure A.1). The different panels in Figures 4 and A.1 show (from top to bottom) the Bry line profile (upper panels), visibilities (middle upper panels), differential phases (middle lower panels), and closure phases (lower panels).

The wavelength-dependent visibilities (middle upper panels of Fig. 4 and A.1) clearly increase across the Bry line-profile at all baselines. This indicates that the Bry emitting region is more compact than the continuum emitting region. Previous spectro-interferometric studies of HD 98922 at medium resolution ($R = 1500$) have also detected an increase in the visibility within the Bry line (Kraus et al. 2008) with baseline lengths up to 60 m. Our HR mode AMBER observations allow us, for the first time, to measure the visibilities and phases in ~ 30 different spectral

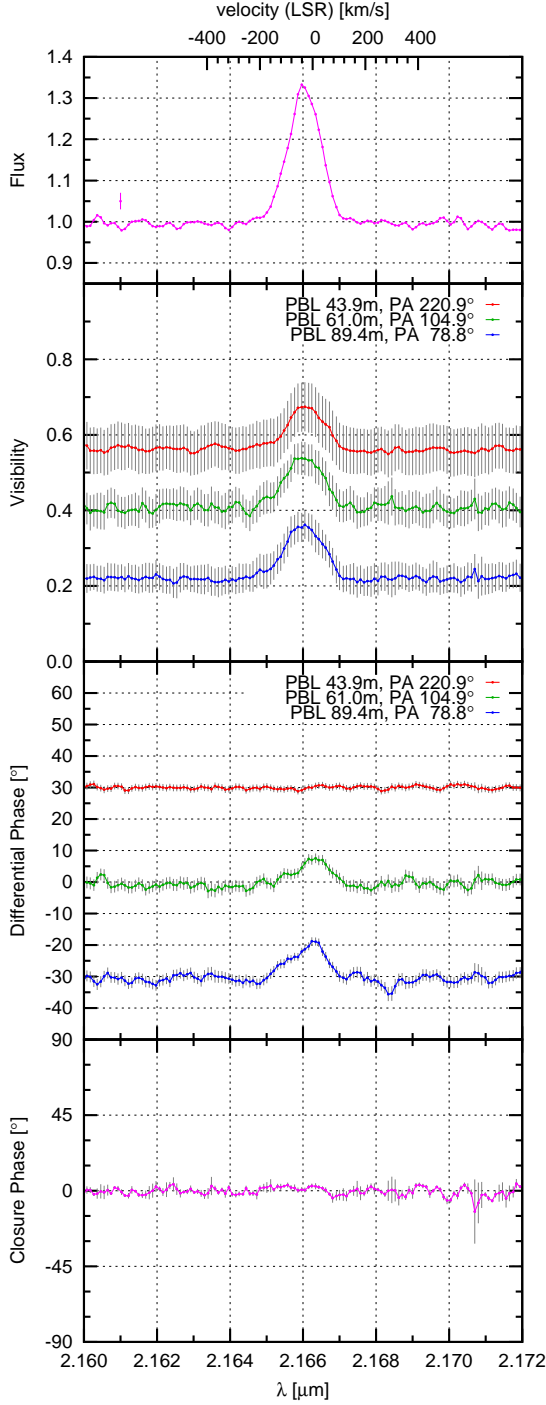


Fig. 4. AMBER observation of HD 98922 at spectral resolution of $R=12\,000$ (28 Feb 2013). From top to bottom: wavelength dependence of flux, visibilities, wavelength-differential phases, and closure phase observed at three different projected baselines (see labels in figure). For clarity, the differential phases of the first and last baselines are shifted by $+30^\circ$ and -30° , respectively.

channels across the Bry line with baseline lengths up to 90 m, providing a higher spatial and spectral resolution. Furthermore, our data indicate that the visibilities from the three baselines are slightly asymmetric with respect to the Bry peak, namely the

visibilities of the blue-shifted wings are systematically smaller than those of the red-shifted wings. This might indicate that, at high velocities, the red-shifted Bry emitting region is slightly more compact than the blue-shifted one, or that there is a highly asymmetric continuum.

Differential phases (Φ ; lower middle panels of Fig. 4 and A.1) measure the displacement of the line-emitting region with respect to the continuum-emitting region. For example, the accuracy reachable with a 90 m baseline and with $\Phi \sim 10^\circ$ is about 0.14 mas, or 0.06 AU at a distance of 440 pc. Therefore the observed displacements provide us with a very sensitive measurement of the gas kinematics on scales of a few stellar radii. For the first time in this object, we detect a clear photocentre shift of the Bry line with respect to its continuum at the 61 m and 89 m baselines. The line displacement is observed in different velocity channels (both blue- and red-shifted). Notably, differential phases in both blue- and red-shifted channels have the same sign, and therefore they do not exhibit the typical “S” shape usually observed in rotating discs. Additionally, the red-shifted channels in both baselines show larger differential phases with respect to the blue-shifted ones. Because these are line-to-continuum phases, the observed shifts in the differential phases could be caused by continuum shifts, line shifts, or both. The maximum value of the line/continuum displacement is 0.02 ± 0.04 mas, 0.14 ± 0.04 mas, and 0.16 ± 0.03 mas for the 43.9 m, 61 m, and 89.4 m baselines, respectively.

Finally, it is also worth noting that the closure phases do not show any significant deviation from zero within the error bars (see lower panels of Figs. 4 and A.1); we measure an average value of $-0.44^\circ \pm 2.55^\circ$.

5.2. Geometric modelling: size of Bry and continuum emitting region

Our AMBER high spectral resolution data also provide us with a measurement of the visibilities for the pure Bry line-emitting region in the different spectral channels across the Bry emission line (see Figure 5). The continuum-corrected line visibilities are fundamental for measuring the size of the Bry line-emitting region. Within the wavelength region of the Bry line emission, the measured visibility has two components: the pure line-emitting part and the continuum-emitting part, which includes continuum emission from both the circumstellar environment and the unresolved central star. Therefore, following Weigelt et al. (2007), the emission line visibility V_{Bry} in each spectral channel can be written as

$$F_{\text{Bry}} V_{\text{Bry}} = \sqrt{|F_{\text{tot}} V_{\text{tot}}|^2 + |F_c V_c|^2 - 2 F_{\text{tot}} V_{\text{tot}} F_c V_c \cdot \cos \Phi}, \quad (2)$$

where F_{Bry} is the wavelength-dependent line flux, V_{tot} (F_{tot}) is the measured total visibility (flux) in the Bry line, V_c (F_c) is the visibility (flux) in the continuum, and Φ is the measured wavelength-differential phase within the Bry line.

In our analysis, we also included the intrinsic Bry photospheric absorption feature of HD 98922, using a synthetic spectrum with the spectral type and surface gravity values as in Table 2.

The size of the continuum was obtained using the AMBER GTO archival data at low resolution. By fitting a circular symmetric Gaussian model, we derived a Gaussian half width at half maximum (HWHM) radius of 1.8 ± 0.3 mas (or $\sim 0.79 \pm 0.13$ AU at a distance of 440 pc). A similar value of 1.6 ± 0.3 mas (or $\sim 0.70 \pm 0.13$ AU) was inferred by fitting a ring model with a ring width of 20% of the inner ring radius.

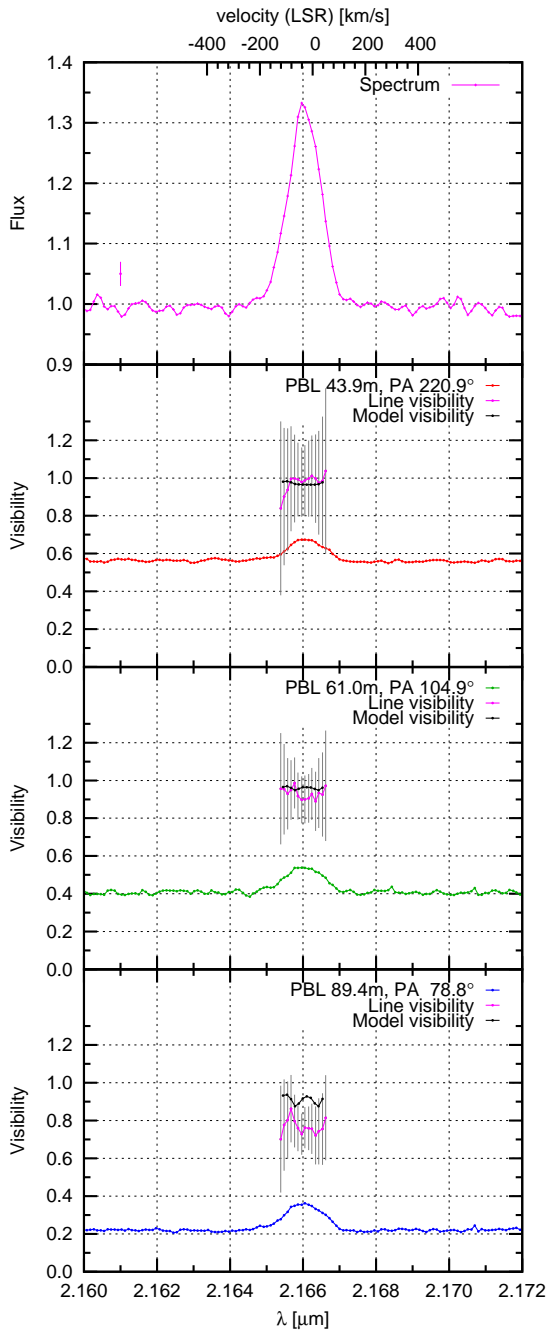


Fig. 5. Comparison of the observed (pink solid lines) and modelled (black solid lines) continuum-corrected (pure) Bry line visibilities of our AMBER observation of HD 98922. From top to bottom: wavelength dependence of flux and visibilities of the first, second, and third baseline. In each visibility panel the observed total visibilities (red, green, blue, as in Fig. 4), the observed pure Bry line visibilities (pink), and the modelled pure Bry line visibilities (black) are shown.

To get reliable line visibility values, we derived the continuum-corrected (hereafter pure) line visibility only in the spectral regions where the line flux is higher than $\sim 10\%$ of the

continuum flux. The results are shown in Figure 5. The average line visibility is $\sim 1 \pm 0.2$ for the shortest baseline (~ 44 m), $\sim 0.9 \pm 0.15$ for the medium baseline (~ 61 m), and $\sim 0.8 \pm 0.1$ for the longest baseline (~ 89 m). For this baseline, the line is spatially resolved even when the error bars are taken into account.

The size of the line-emitting region was obtained by fitting a circular symmetric Gaussian model to the line visibilities presented in Fig. 5. We obtain a Gaussian HWHM radius of 0.65 ± 0.09 mas or $\sim 0.29 \pm 0.04$ AU at a distance of 440 pc. Similarly, by fitting a ring model with a ring width of 20% of the inner ring radius, an inner radius of 0.70 ± 0.09 mas, or $\sim 0.31 \pm 0.04$ AU, is inferred.

6. Disc and disc-wind models

To constrain the physical processes taking place in the innermost disc regions, a proper physical modelling is required. As in our previous interferometric works (see Weigelt et al. 2011; Garcia Lopez et al. 2015), to analyse the observed Bry line profile, line visibilities, differential phases, and closure phase we employed a previously developed disc-wind (DW) model (see Sect. 6.1) and developed a continuum disc model; both models are adapted to the HD 98922 stellar parameters and observations.

6.1. Disc-wind model

A detailed description of our DW model, its parameters, and the algorithms used for the model computation can be found in our previous papers (see Grinin & Tambovtseva 2011; Weigelt et al. 2011; Tambovtseva et al. 2014; Garcia Lopez et al. 2015). In the following, we briefly outline the model's main characteristics and define its free parameters, which are listed in Table 3.

The developed model is a warm disc-wind model (see e.g. Safier 1993; Garcia et al. 2001), including only hydrogen atoms with constant temperature ($\sim 10\,000$ K). The wind is rapidly heated by ambipolar diffusion to a temperature of $\sim 10\,000$ K, and the wind electron-temperature (T_e) in the acceleration zone near the disc surface is not high enough to excite the Bry line emission, which is emitted further out along the wind streamlines.

Briefly, the disc wind is launched from a disc inner radius ω_1 to an outer radius ω_N (called wind footpoints), where N is the number of streamlines. Its half opening angle (θ) is defined as the angle between the innermost wind streamline and the system axis. The local mass-loss rate per unit area on the disc surface is defined as $\dot{m}_w(\omega) \sim \omega^{-\gamma}$, where γ is the mass-loading parameter that controls the ejection efficiency. Therefore, the total mass-loss rate (\dot{M}_w) is

$$\dot{M}_w = 2 \int_{\omega_1}^{\omega_N} \dot{m}_w(\omega) 2\pi \omega d\omega. \quad (3)$$

Finally, the β parameter in Table 3 regulates the acceleration of the wind along the streamlines (see Tambovtseva et al. 2014). Some model parameters such as \dot{M}_w , the stellar parameters, and the disc inclination angle were derived in this work or were taken from the literature (see Tab. 2). The explored range of values of \dot{M}_w varies from ~ 0.05 to 0.3 times the inferred mass accretion rate (\dot{M}_{acc} , see Table 2). The value for the corotation radius (R_{cor}^1) obtained from the $v_{rot} \sin i$ value (Table 2) is

¹ with $R_{cor} = (GM_*/v_\omega^2)^{1/3}$, where G is the gravitational constant, M_* the stellar mass, and v_ω the stellar angular velocity.

$R_{cor}=1.55 R_*$. The estimated value of R_{cor} is then set as a lower limit to the inner disc-wind launching radius value ω_1 to allow for the presence of a hypothetical magnetosphere within this region.

To find the best model that reproduces our observations, we adopt the following steps (for more details, see also Weigelt et al. 2011; Grinin & Tambovtseva 2011). First, we fit the observed Bry line profile and intensity. To this end, we solve the equations of statistical equilibrium and compute the population of the H I atomic levels and the ionisation degree along each streamline using the numerical codes developed by Grinin & Mitskevich (1990) and Tambovtseva et al. (2001) for moving media. Second, we calculate the intensity of the line radiation in the entire emitting volume filled in by the disc wind and compare both modelled and observed line profiles. When a good agreement is found, we compute the brightness distribution of the disc wind and the disc continuum at different radial velocities, i.e. a two-dimensional intensity distribution map, which provides us with the interferometric observables. The modelling results are, in this way, directly compared with the observations until the best-fit model is found. The best model is that which best matches, within the error bars, all eleven interferometric observables (line profile shape and intensity, visibilities, pure line visibilities, differential phases, and closure phase) of our observation within the selected range of parameters listed in Table 3.

This latter reports the free model parameters, the explored range of parameters (Column 2), and the values of our best-fit model, called P5 (Column 3 of Tab. 3). P5 model was selected after checking hundreds of different models configurations (see Column 2, Tab. 3).

6.2. Continuum disc model

To reproduce our AMBER HR interferometric observations, we have assumed that the HD 98922 continuum emission consists of three different components: the star, an inner disc, and an outer dusty disc (see e.g. Tannirkulam et al. 2008b; Dullemond & Monnier 2010).

The main constraints for our developed temperature-gradient model of the continuum emission are provided by: *i*) the $F_{obs}(K)/F_*(K)$ ratio, namely the ratio between the HD 98922 observed flux continuum in the K -band (star plus accretion disc) and the theoretical stellar flux, and *ii*) the absolute visibilities of the continuum. The stellar continuum at the Bry rest wavelength ($F_*(K)$) was computed from a Kurucz synthetic spectrum with $T_{eff}=10400$ K and $\log g=3.5$ (see Sect. 3.1.1 and Tab. 2), whereas the $F_{obs}(K)$ is computed from the YSO K -band magnitude (see Tab. 2). The adopted disc inclination angle is the same as for the disc-wind model ($i=20^\circ$, see Tab. 2). To fit the disc continuum level of our observations, we use an inner gaseous disc plus an outer dusty disc, as in Tannirkulam et al. (2008a) and Benisty et al. (2010). The sum of fluxes from the gaseous and dusty disc matches the observed flux close to the Bry wavelength. The inner disc has constant brightness distribution and ranges from $R_m=4 R_*$ (or 0.14 AU) to $R_S=20 R_*$ (or 0.7 ± 0.2 AU). The dust temperature (T_{d0}) at this radius (R_S) is 1500 K, that is, the silicate dust sublimation temperature (see e.g. Dullemond & Monnier 2010). The dust temperature exponentially decreases with the distance r from the star accordingly with the following power law: $T_d(r) = T_{d0}(r/R_S)^\alpha$, with $\alpha=-0.5$. The intensity of the radiation in the continuum was calculated in the blackbody approximation. The disc brightness is distributed symmetrically along azimuthal direction. Fig. 6 (upper panel)

shows the disc continuum model plus the Bry intensity distribution map (for $v_r=0$ km s $^{-1}$) of the best-fit disc-wind model (P5).

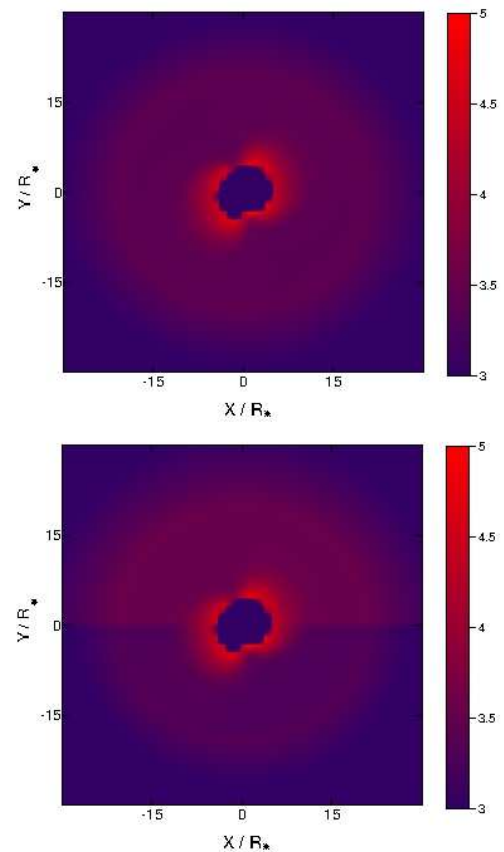


Fig. 6. *Upper panel:* Bry intensity distribution map (for $v_r=0$ km s $^{-1}$) of the best-fit disc-wind model (P5) overlapped on the map of the disc continuum with symmetric brightness distribution. X and Y coordinates are in R_* . *Lower panel:* Same as in the upper panel, but for the disc continuum with asymmetric brightness distribution.

6.3. Modelling results

As seen in Sect. 5, the Bry emission line shows V-shaped differential phases and slightly asymmetric visibilities (see Fig. 4).

Because of the aforementioned shape of the differential phases, our first attempts at modelling these interferometric observables with our disc-wind model plus the continuum disc model described in Sect. 6.2 were unsuccessful (see Fig. B.1 in Appendix B). Indeed, the combination of these disc and disc-wind models produces, as expected, symmetric S-shaped differential phases (originating from disc rotation, see Figure 7, upper panel) and symmetric visibilities in the Bry line.

In principle, the unusual shapes of the observed visibilities and differential phases, which are a measurement of the photocentre shifts, can be caused by an asymmetric Bry line-emitting region, an asymmetric continuum-emitting region, or by a combination of both. As a consequence, the differential phase in the blue or red wings of the Bry can be displaced, and the visibilities become asymmetric. An asymmetric Bry line-emitting region could be a plausible explanation for the observed differential phase. In theory, this might originate from a collimated wind/jet, whose red-shifted lobe is hidden by the disc (see e.g. Benisty et al. 2010). However, in this case, we would observe a blue-shifted photocentre shift, whereas our observations clearly

Table 3. Disc-wind accretion model parameters: range of explored values and values for our best-fit model P5

Parameter ^a	Range of Values	Model P5
Half opening angle (θ_1)	20°–60°	30°
Inner radius - $\omega_1(R_*)$	2–3 (0.07–0.1 AU)	3 (0.1 AU)
Outer radius - $\omega_N(R_*)$	4–30 (0.13–1 AU)	30 (1 AU)
Acceleration parameter (β)	3–7	5
Mass load parameter (γ)	2–5	3
Mass-loss rate - $\dot{M}_w(M_\odot \text{ yr}^{-1})$	10^{-8} – 3×10^{-7}	2×10^{-7}

Notes. ^(a) See Weigelt et al. (2011) for definition and detailed description of the listed disc-wind model parameters

show a red-shifted peak in the differential phases. On the other hand, it is worth noting that the HD 98922 NIR continuum emission shows an asymmetric brightness distribution on small and large spatial scales (from ~ 1 AU to tens of AUs), as observed in both VLT/PIONIER and SINFONI data (see Sect. 4, and Kluska et al. 2014). Therefore, an asymmetric continuum might be responsible for the observed shape in the differential phases and visibilities.

To verify this hypothesis, we modified the disc model of Sect. 6.2, developing a very simple asymmetric continuum disc model, which mimics the difference in brightness between the two halves of the disc. However, it is beyond the scope of this paper to model the observed disc asymmetries in details and at small scales. Therefore, for the sake of simplicity, the intensity of the radiation was increased by a factor f_c in the northern half of the disc and decreased by the same factor in the southern half of the disc of both gaseous and dusty disc regions. As a consequence, the disc brightness from PA 270° to 90° (i.e. the northern half of the disc) is brighter than the southern half of the disc (extending from PA 90° to 270°). Lower panel of Figure. 6 shows the asymmetric disc continuum model plus the Br γ intensity distribution map (for $v_r=0 \text{ km s}^{-1}$) of the best-fit disc-wind model (P5). The brightness contrast (f_c) between the two halves of the disc is a free parameter of the model. The asymmetric disc was then rotated by 10° steps until we obtained a solution that best fits both visibilities and differential phases at the three observed baselines. In order to fit the continuum level of the three baselines, we changed the f_c value, which was varied from 1 to 4. The best continuum model has an f_c value of 2.5 and a disc rotation of 70°. Although very simple, our asymmetric disc model, combined with the disc-wind model described in Sect. 6.1, is able to reproduce the V-shaped differential phases (see Fig. 7, middle panel). Middle and lower panels of Fig. 7 present differential phases for different rotation of the disc and the disc wind: counter clockwise (b ; middle panel) and clockwise rotation (c ; lower panel). It should be noted that the shapes of the differential phases in the two cases are slightly different. The reason is the disc wind. By modifying the direction of the disc-wind rotation, the distribution of the disc-wind brightness changes but the distribution of the disc brightness remains unchanged. Thus, the differential phases are sensitive to the direction of the disc and disc-wind rotation.

Figure 8 shows the Br γ intensity distribution maps of our best disc-wind model (P5) in logarithmic scale, as a function of different Br γ radial velocities. The adopted system axis inclination angle with respect to the line of sight is 20° (see Table 2; Hales et al. 2014). Column 3 of Table 3 reports the values of the model free parameters. Figure 9 shows a comparison of our best model with the interferometric observables. Despite the simplicity of the proposed disc model, the interferometric ob-

servables are approximatively well reproduced, demonstrating that V-shaped differential phases and slightly asymmetric line visibilities may originate from the uneven brightness distribution of the disc. It is also worth to note that the value derived for f_c (i.e. 2.5) is, within the error bars, identical to the average brightness ratio between the northeastern and southwestern disc regions of the SINFONI spectral images (see Sect. 4), and that the PA of our asymmetric disc model (i.e. 70°) is compatible with the observed asymmetric disc orientation in Fig. 3. In Fig. 5, we compare observational and model results of the Br γ pure line visibilities (pink and black solid lines, respectively). Our model approximatively reproduces the observed values at the shortest and medium baselines (upper and central panels), but it slightly overestimates the pure line visibility at the longest baseline (lower panel). This result represents a good compromise. To obtain a lower value of the modelled pure line visibility at the longest baseline, it would be necessary to increase the size of the disc-wind region (i.e., the value ω_N). On the other hand this would increase the values of the differential phase, and the resulting model would not match our observations.

Moreover, we also note that the Br γ emitting region is spatially extended. As a consequence, it cannot be modelled with a compact, spatially unresolved component alone, such as a magnetosphere (size $< 2\text{--}3 R_*$; see e.g. Tambovtseva et al. 2014; Garcia Lopez et al. 2015), that cannot be spatially resolved at these baselines. Therefore we conclude that the Br γ emission must be at least a combination of a compact (i.e. spatially unresolved) emission (magnetosphere) plus an extended spatially resolved emission (i.e. the disc wind). The addition of a compact region (i.e. visibility value = 1 at all baselines) implies that the size of the spatially resolved component must be enlarged to match the observed visibility values. This can be achieved with a larger disc-wind footpoint radius ω_N , which, in turn, gives larger values of the differential phases. In addition, changing the footpoint radius modifies the line profile as well. Thus our model has to take into account these three observables (visibility, differential phase, and line profile). In conclusion, by comparing our model with the observations, we do not rule out that some of the Br γ emission may originate from the magnetosphere (i.e. a spatially unresolved region), however we can exclude that this latter is the main mechanism, which produces the Br γ emission (see also Garcia Lopez et al. 2015).

Finally, it is worth comparing the mass-loss rate derived from our model with the mass accretion rate inferred from the observations, as their ratio (\dot{M}_w/\dot{M}_{acc}) usually provides constraints on the efficiency of the accretion/ejection mechanism in young stars, typically ranging from 0.1 to 0.3 (see e.g. Shu et al. 1988; Ferreira et al. 2006). In our case, the accuracy of the \dot{M}_{acc} value is $\sim 30\%$. This uncertainty mainly originates from the employed empirical relationship (see e.g. Mendigutía et al. 2011, 2013).

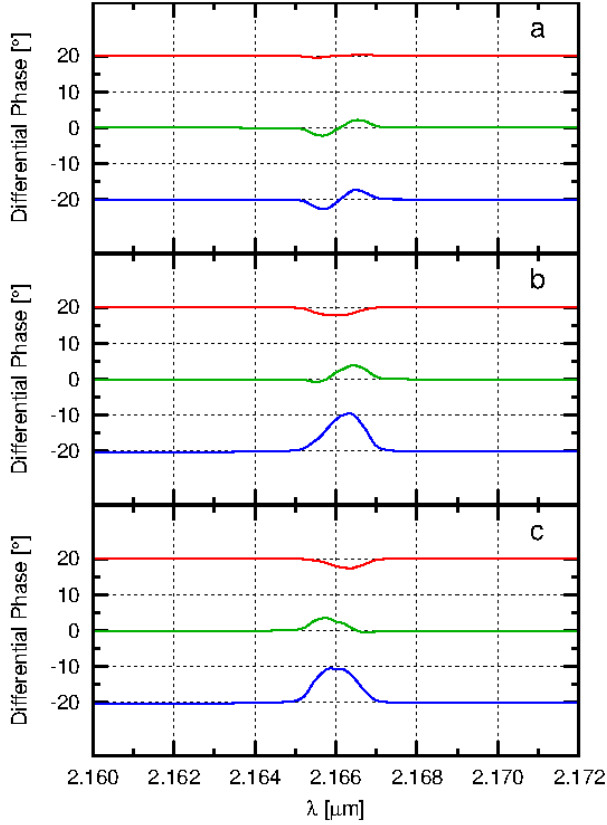


Fig. 7. Differential phases for the three observed baselines (as in Fig. 4), calculated for the best disc-wind model P5 plus: symmetric disc model (a, upper panel), asymmetric disc model with a counter clockwise (b, middle panel), and clockwise (c, lower panel) disc and disc-wind rotation. For clarity, the differential phases of the first and last baselines are shifted by $+20^\circ$ and -20° , respectively.

On the other hand, it is more difficult to estimate the accuracy of the \dot{M}_w value as inferred from our best-fit model, but it is probably of the same order as the \dot{M}_{acc} value, owing to the use of our simplified model of the disc wind and to the uncertainties in the fit of the Bry line profile and intensity. As a result, their ratio is well within the expected range of values. It should also be noted that in our best-fit model of HD 98922 the disc wind starts in the close vicinity of the star. This can mean that the wind at least partially originates as a result of interaction of the accretion disc with the magnetosphere. In this case the \dot{M}_w/\dot{M}_{acc} ratio can be higher than that used in the theory of the magneto-centrifugal disc wind (0.1), for example due to the effect of the mechanism of the magnetic propeller (see e.g. Illarionov & Sunyaev 1975).

7. Discussion

7.1. An asymmetric disc emission

Asymmetries in the disc brightness distribution are not uncommon as they have been observed at NIR wavelengths in several Herbig Ae/Be stars and CTTSs with the Hubble Space Telescope (HST) or with ground-based telescopes through AO assisted imaging or interferometric imaging (see e.g. Krist et al. 2005; Dullemond & Monnier 2010). These images of spatially resolved discs display a wide range of spectacular asymmetries, including arcs (e.g. Fukagawa et al. 2006; Krist et al. 2012), gaps and dips (e.g. Krist et al. 2005), warps (e.g. Golimowski et al.

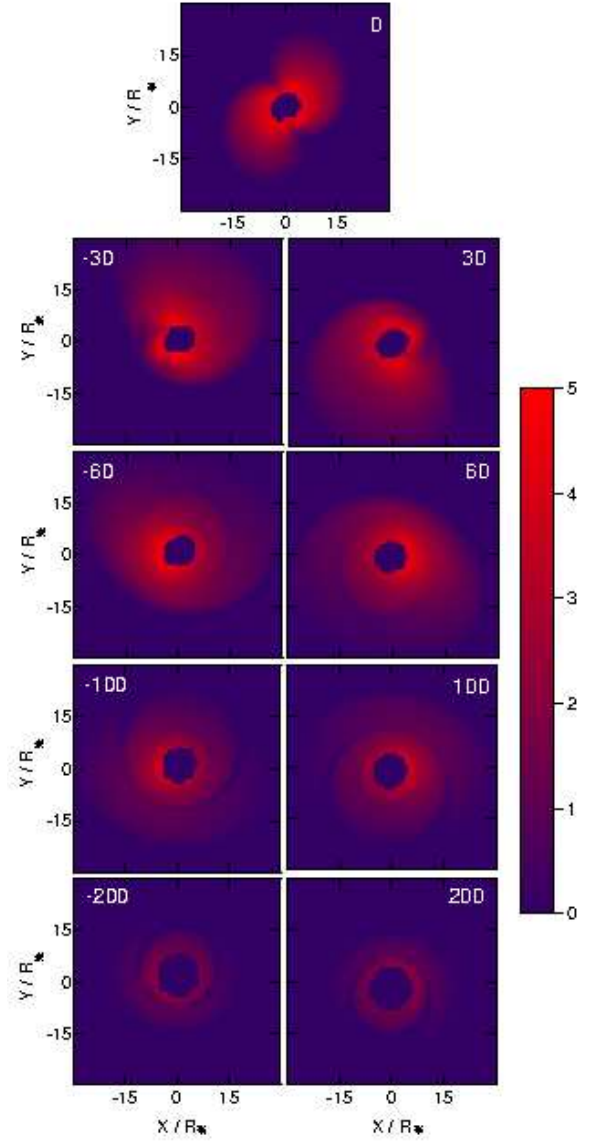


Fig. 8. Bry intensity distribution maps of our best disc-wind model P5 (see values of parameters in Table 3, Column 3) in logarithmic scale (arbitrary units). Each panel shows the intensity map for a different radial velocity, indicated by a white label in units of km s^{-1} . X and Y coordinates on each map are in R_* . The system axis inclination angle with respect to the line of sight is assumed to be 20° (see Table 2). Continuum emission from the disc and central star, located at coordinates 0,0, are not shown.

2006; Marino et al. 2015), or more complex morphologies (e.g. Hines et al. 2007; Kalas et al. 2007; Mazoyer et al. 2014).

The origin of such asymmetries in discs varies from case to case, but mostly involves disc geometry, illumination (e.g. only one side of the illuminated inner rim is observed; see Monnier et al. 2006), dust grain composition and distribution, or tidal effects: *a*) brightness asymmetries between the two halves of the disc may originate from different emission in the forwards/backwards scattered light in inclined discs (see e.g. Krist et al. 2005; Dullemond & Monnier 2010); *b*) gaps and dips in discs can be produced by shadowing effects caused by optically thick clouds of dust and gas (see e.g. Krist et al. 2005); *c*) arcs of different brightness and size may originate from shadowing effects, produced by a different inclination between the inner

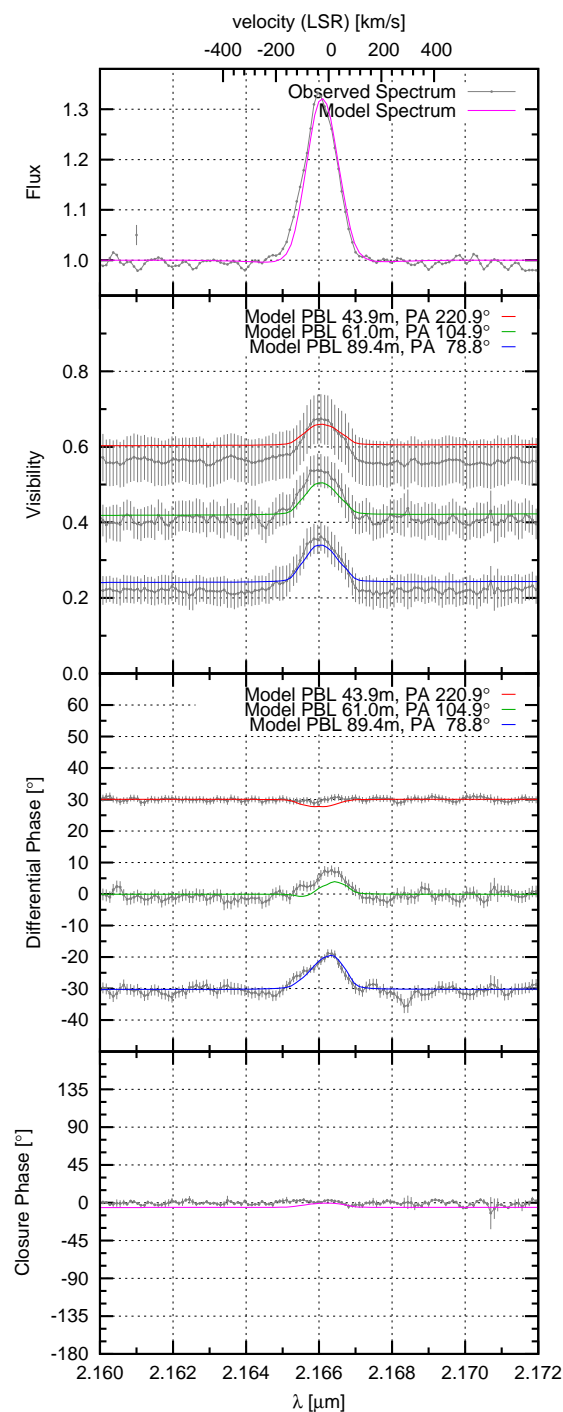


Fig. 9. Comparison of our interferometric observations with the interferometric observables derived from our best disc-wind model P5 plus asymmetric disc model. From top to bottom, observed Br γ line profile (grey) and model line profile (pink), observed visibilities (grey dots with error bars) and model visibilities (coloured lines), observed and model wavelength-differential phases, and observed and model closed phase.

and the outer disc (see e.g. Marino et al. 2015); *d*) arcs can be also produced by dust density variations and stellocentric grain-size segregation; *e*) arcs and rings can also originate from warped

discs produced by tidal interactions (see e.g. Tuthill et al. 2001; Demidova & Grinin 2014; Marino et al. 2015).

To investigate the origin of the asymmetries in the HD 98922 disc intensity distribution, we should analyse large- and small-scale asymmetries.

Asymmetries on small scales include *i*) asymmetric inner brightness distribution, due to shadowing caused by the vertically extended structure of the inner rims (which can appear unusually asymmetric if the disc geometry is not exactly face-on; see e.g. Monnier et al. 2006), or *ii*) brightness peaks and gaps along the disc azimuthal angle.

The large scale asymmetry is represented by the different brightness distribution between the NNE and SSW side of the disc (whose ratio is about 2.5). This feature is detected at both VLTI (AMBER and PIONIER) and SINFONI spatial scales.

The different emission between forward and backward scattering could be the simplest explanation. The fact that the brightness asymmetry lies approximately in the direction of the system axis also supports this hypothesis. On the other hand, this asymmetry is observed both at large (SINFONI) and small (VLTI) scales, and the SINFONI data are likely dominated by scattered light, whereas the VLTI data are possibly dominated by thermal emission. These considerations are in favour of an intrinsic perturbation of the disc.

In principle, disc warping is the most plausible explanation. Recently Demidova & Grinin (2014) modelled the disc asymmetry of LkHa 101 (a Herbig Ae star; see e.g. Tuthill et al. 2001), which shows a warped disc. Such asymmetry is very similar to that observed in HD98922 (see e.g. Figure 1 in Tuthill et al. 2001). The warp is generated by the tidal effects of a low-mass companion (with mass ratio q ranging from 0.1 to 0.01) with its orbit slightly inclined with respect to the disc plane. The warp height varies as the companion extracts matter out of the mid-plane of the disc. These perturbations generate an azimuthal asymmetry of the extinction. As a result, the illumination of the disc by the star becomes asymmetric.

In the case of HD 98922, the SINFONI data exclude the presence of a companion as close as ~ 20 AU from the central source down to $\sim 0.5 M_{\odot}$ (see Sect. 4, i.e. $q \geq 0.1$). Moreover, the VLTI/PIONIER reconstructed image (see Kluska et al. 2014) does not show any evidence of a close binary (up to a few AUs from the star), but for these data no sensitivity limit is provided. Indeed, given the VLTI/PIONIER sensitivities (see e.g. Absil et al. 2011), the dynamical range of the reconstructed image should be lower than or equal to 100 (Kluska priv. comm). This roughly gives a lower limit of $H \sim 10.2$ mag or $M_H \sim 2$ mag at the adopted distance of 440 pc. By adopting the Siess et al. (2000) evolutionary tracks (see Sect. 4), we obtain an upper limit for the mass of any undetected companion of ~ 0.4 – $0.5 M_{\odot}$. Our reasoning implies that the presence of a low-mass object ($M_* \leq 0.5 M_{\odot}$) or a massive protoplanet cannot be ruled out by both SINFONI or VLTI observations. As a consequence the hypothesis of a warped disc, originating from a very low-mass companion (i.e. with mass ratio in the range of 0.1–0.01), still holds.

7.2. Inner circumstellar structure of HD 98922

Our interferometric observations show that the Br γ line visibilities are higher than those of the continuum, indicating that the Br γ emitting region is less extended (~ 0.3 AU in radius) than the continuum emitting region (~ 0.7 AU in radius). The pure Br γ line visibility of HD 98922 is 0.8 at the longest baseline (89 m), indicating that the Br γ originates from a spatially ex-

tended region. Our interferometric observables (Bry line profile, visibilities, differential phases, closure phases) are sufficiently well reproduced by a disc-wind model with a half opening angle of 30° , wind footpoints extending from ~ 0.1 AU to ~ 1 AU, and $\dot{M}_w \sim 2 \times 10^{-7} M_\odot \text{ yr}^{-1}$.

The wind mainly originates from the inner disc region, beyond the corotation radius (R_{cor}), which can be placed at ~ 0.06 AU from the star (assuming $v_{rot} \sin i = 50 \pm 3 \text{ km s}^{-1}$ and $i = 20^\circ$, see Table 2). Notably, the wind model extends slightly beyond the dust sublimation radius, located at ~ 0.7 AU. This is required to fit both the low visibilities and the large differential phases. Indeed, a more compact disc-wind model would produce small or no differential phases and it would not reproduce our interferometric data (see Garcia Lopez et al. 2015).

It is worth mentioning that some of the previous HR AMBER/VLTI studies of Herbig AeBe stars (e.g. MWC 297 and HD 163296: see Weigelt et al. 2011; Garcia Lopez et al. 2015, respectively) show similar results to those described here, namely that the Bry line is emitted from an extended region that can be modelled with a disc wind. However, MWC 297 (SpT=B1V) and HD 163296 (SpT=A1V) show disc-wind geometries that differ from the one modelled in HD 98922 (SpT=B9V). In the first case, the disc wind was entirely launched from a region located in the continuum disc, i.e. well beyond the inner radius of the continuum-emitting disc (Weigelt et al. 2011). The disc-wind model of MWC 297 extends from ~ 17 to $35 R_*$ with a large half opening angle of 80° . Moreover, the values of the β and γ parameters, which regulate the mass load along the streamlines, are quite small (1 and 2, respectively). This indicates that the mass load is shifted towards the outer streamlines of the wind. Weigelt et al. (2011) suggest that such features might be caused by the intense radiation pressure of the central source. On the other hand, HD 163296 has a very compact disc wind (extending ~ 0.16 AU), confined to the inner gaseous disc (Garcia Lopez et al. 2015). In this case, ω_1 and ω_N radii are 2 and $4 R_*$ with a half opening angle of 45° . In contrast to MWC 297 and HD 98922, HD 163296 drives a well-collimated jet (see e.g. Wassell et al. 2006); therefore, the smaller ω values might be related to this. The modelled disc wind of HD 98922 has an intermediate geometry, with the Bry line-emitting region located mostly in the inner disc, but extending slightly beyond the dust sublimation radius. Our disc-wind model extends from 3 to $30 R_*$ with a half opening angle of 30° . Although spectroscopy indicates a strong outflow originates from its disc (e.g. this work; see also Grady et al. 1996), no signature of any collimated jet has been detected so far. On the other hand, the β and γ parameters are identical to those of HD 163296, suggesting that the mass load is similarly distributed along the stream lines, but along a larger portion of the disc.

Following these considerations, our modelling seems to suggest that different types of outflows/jets can be modelled with different disc-wind geometries. Moreover, these disc-wind geometries seem to display an evolutionary trend, following the different spectral types and/or evolutionary stages of the central sources. It is also tempting to speculate on some additional correlations between, for example, the Bry size, its location, and the stellar type and mass of the central source. If our previous speculations are correct, we would then expect that disc wind geometry, kinematics, and Bry size depend on the central source, namely its mass and/or evolutionary stage. Nevertheless, it is also clear that the sample so far studied is still too limited to draw any firm conclusion.

8. Conclusions

In this paper, we analyse the main physical parameters and the circumstellar environment of the young Herbig Be star HD 98922. We present AMBER high spectral resolution ($R=12\,000$) interferometric observations across the Bry line, as well as UVES (high-resolution UVB spectroscopy) and SINFONI-AO assisted (NIR integral field spectroscopy) ancillary data. To interpret our observations, we also developed a magneto-centrifugally driven disc-wind model along with a very simple asymmetric continuum disc model. The main results of this work are the following:

- UVES high-resolution spectrum displays H I photospheric absorption lines from the Balmer series (from H 11 to H β) with superimposed emission originating from circumstellar activity. Additional weaker photospheric absorption features from atomic lines (He I, Ca II, Fe I, Fe II, Mg I, Mg II, Si I, Si II) and circumstellar emission lines (Fe II and Ti II) are also detected in the spectrum.
- Our analysis of the UVES spectrum indicates that HD 98922 is a young (5×10^5 yr) Herbig Be star (SpT=B9V), located at a distance of $440 \pm_{50}^{60}$ pc, with $\dot{M}_{acc} = (9 \pm 3) \times 10^{-7} M_\odot \text{ yr}^{-1}$.
- SINFONI K-band AO-assisted imaging shows a spatially resolved circumstellar disc (~ 140 AU in diameter) with asymmetric brightness distribution. In particular the NNE side is ~ 2.5 times brighter than the SSW side. Other features at smaller spatial scales include peaks and dips along the disc.
- VLTI/AMBER high spectral resolution observations ($R=12\,000$) allow us to study the interferometric observables (i.e. line profile, visibilities, differential phases, and closure phases) in many spectral channels across the Bry line. Differential phases are V-shaped and line visibilities are slightly asymmetric.
- Our interferometric observations indicate that the Bry line visibilities are higher than those of the continuum, i.e. the Bry emitting region is less extended than the continuum emitting region. For the first time, HD 98922 continuum-corrected Bry pure line visibilities are 0.8 at the longest baseline (89 m), indicating that the Bry originates from a spatially extended region. By fitting geometric Gaussian and ring models to the derived pure line visibilities, we infer a HWHM Gaussian radius of 0.65 ± 0.09 mas (or $\sim 0.29 \pm 0.04$ AU at a distance of 440 pc) and ring-fit radius of 0.70 ± 0.09 mas ($\sim 0.31 \pm 0.04$ AU) for the Bry emitting region. The Bry line radius is smaller than the inner rim radius of 0.7 AU of our continuum temperature model.
- To obtain a more physical interpretation of our Bry AMBER observations, we employed our own line radiative transfer disc-wind model (see Weigelt et al. 2011; Grinin & Tambovtseva 2011; Tambovtseva et al. 2014; Garcia Lopez et al. 2015). We computed a model that approximately reproduces all the interferometric observables. As a result, our modelling suggests that the observed Bry line-emitting region mainly originates from a disc wind with a half opening angle of 30° , wind footpoints extending from ~ 0.1 AU to ~ 1 AU, and $\dot{M}_w = 2 \times 10^{-7} M_\odot \text{ yr}^{-1}$.
- The observed V-shaped differential phases and slightly asymmetric visibilities can be reproduced by combining a simple asymmetric continuum disc model with our Bry disc-wind model.

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Appendix A: AMBER interferometric results 26
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Appendix B: Modelling results for disc-wind model
P5 and symmetric disc model

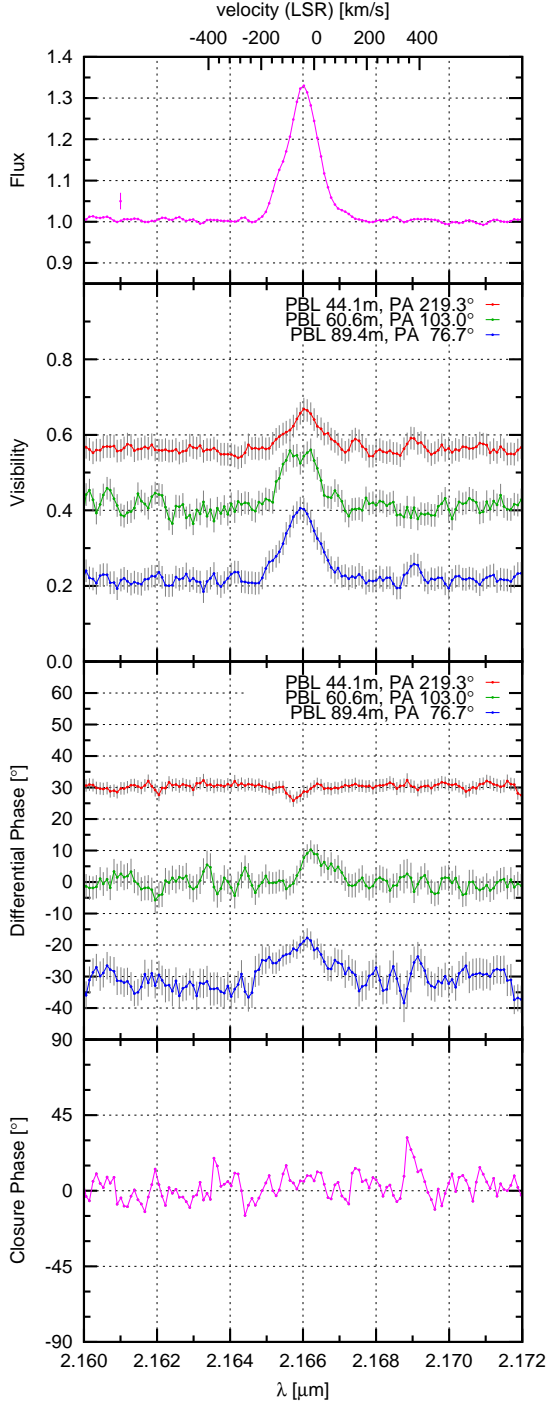


Fig. A.1. AMBER observation of HD 98922 at spectral resolution of $R=12\,000$ (26 Dec 2012). From top to bottom: wavelength dependent flux, visibilities, wavelength-differential phases, and closure phases observed at three different projected baselines (see labels in figure). For clarity, the differential phases of the first and last baselines are shifted by $+30^\circ$ and -30° , respectively.

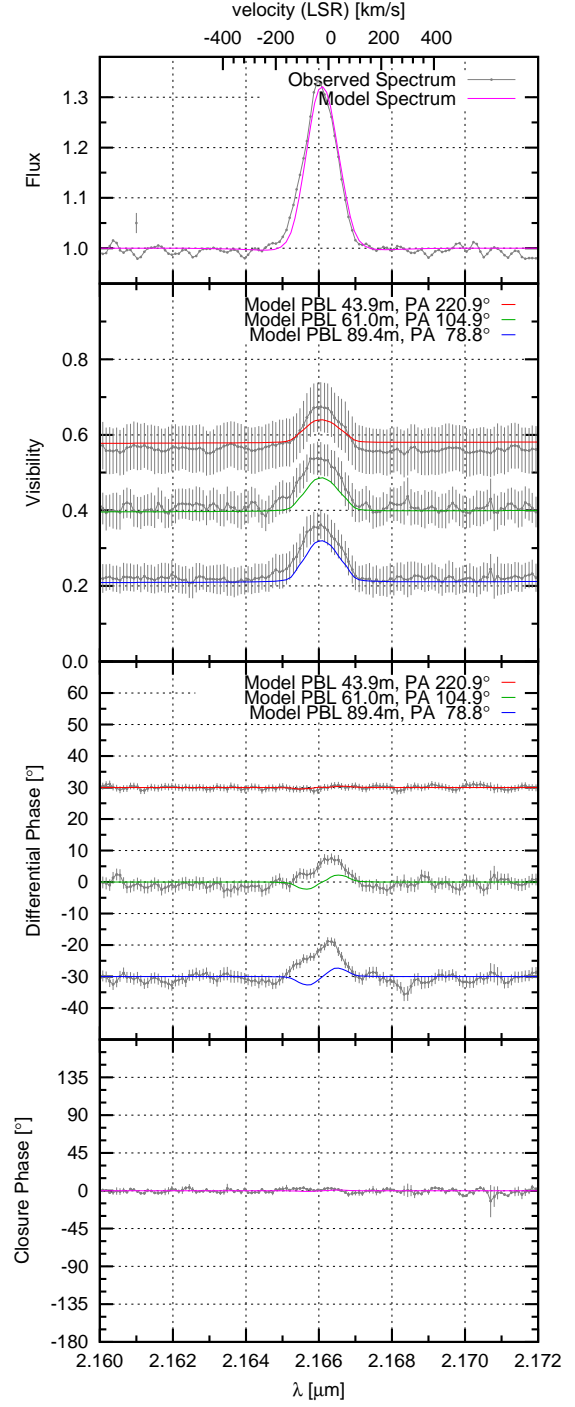


Fig. B.1. Comparison of our interferometric observations with the interferometric observables derived from our best disc-wind model P5 plus the symmetric disc model described in Sect. 6.2. From top to bottom, observed Br γ line profile (grey) and model line profile (pink), observed visibilities (grey dots with error bars) and model visibilities (coloured lines), observed and model wavelength-differential phases, and observed and model closure phase.